

# Characterizing the properties of cluster precursors in the MALT90 survey

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## ABSTRACT

In the Milky Way there are thousands of stellar clusters each harbouring from a hundred to a million stars. Although clusters are common, the initial conditions of cluster formation are still not well understood. To determine the processes involved in the formation and evolution of clusters it is key to determine the global properties of cluster-forming clumps in their earliest stages of evolution. Here, we present the physical properties of 1244 clumps identified from the MALT90 survey. Using the dust temperature of the clumps as a proxy for evolution we determined how the clump properties change at different evolutionary stages. We find that less-evolved clumps exhibiting dust temperatures lower than 20 K have higher densities and are more gravitationally bound than more-evolved clumps with higher dust temperatures. We also identified a sample of clumps in a very early stage of evolution, thus potential candidates for high-mass star-forming clumps. Only one clump in our sample has physical properties consistent with a young massive cluster progenitor, reinforcing the fact that massive protoclusters are very rare in the Galaxy.

**Key words:** surveys – stars: formation – ISM: clouds – submillimetre: ISM.

## 1 INTRODUCTION

Stars form within dense ( $>10^4 \text{ cm}^{-3}$ ) small cores ( $<0.1 \text{ pc}$ ). Most stars are not born in isolation: these small cores are usually embedded in larger structures commonly referred as clumps ( $\sim 1\text{--}2 \text{ pc}$ ) that have a wide range of masses, sizes and are located in different environments. These clumps will evolve to harbour stars and to form complex stellar systems, or clusters. In the Milky Way there are thousands of open clusters, each harbouring from 100 to  $10^4$  stars, and hundreds of globular clusters, each harbouring from  $10^4$  to over  $10^6$  stars (e.g. Bressert et al. 2012). However, although clusters are common in the Galaxy, the initial conditions of cluster formation are still not well understood.

To understand how clusters are formed it is necessary to identify and characterize the physical properties of cluster-forming clumps in a range of evolutionary stages: from starless clumps to those harbouring deeply embedded young stellar objects.

Because of the relative proximity of low-mass star-forming regions, clumps giving rise to low-mass stars have been systematically

better characterized than their intermediate- and high-mass counterparts (e.g. Tafalla et al. 2004; André et al. 2007). High-mass star-forming clumps, however, are usually located at greater distances and embedded in complex environments which have made their study more challenging. To overcome these difficulties, recent unbiased Galactic plane surveys have been key to determining the physical properties of cluster-forming clumps and to identifying protocluster candidates (e.g. Ginsburg et al. 2012; Hoq et al. 2013; Jackson et al. 2013; Urquhart et al. 2014; Guzmán et al. 2015; Svoboda et al. 2016).

To determine if clumps can evolve into clusters harbouring high-mass stars it is necessary to find the physical properties of the cluster-forming clump progenitors. Kauffmann et al. (2010b) used observations of local molecular clouds to determine an empirical threshold for clouds harbouring high-mass stars. They showed that these clouds follow an empirical relationship between their size and mass, defining a threshold for high-mass star-forming clouds given by  $M(r) \geq 870 M_{\odot} (r/\text{pc})^{1.33}$ , with  $r$  the radius of the cloud.

In a similar manner, one can determine the potential of the clumps to evolve into open or young massive clusters (YMCs). Portegies Zwart, McMillan & Gieles (2010) determined the global properties of open clusters in the Galaxy, providing constraints

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to the physical properties that open cluster progenitors ought to have. YMCs, with masses of  $\geq 10^4 M_\odot$ , might represent the missing link between open ( $M < 10^4 M_\odot$ ) and globular clusters ( $M \sim 10^4\text{--}10^6 M_\odot$ ). Therefore, the study of the physical properties of the YMC progenitors may provide key information about the formation mechanisms of globular clusters, and might contribute to our understanding of cluster formation in the early Universe (Longmore et al. 2014). Clumps forming YMCs ought to have masses  $\geq 3 \times 10^4 M_\odot$  (assuming a star-forming efficiency of 30 per cent). Their gravitational escape velocity needs to be greater than the photoionized gas sound speed so that ionization of newly formed stars will not disrupt the natal gas and the cluster will remain bound. Since observations of YMCs suggest that they are in – or close – to virial equilibrium at ages of  $\sim 1$  Gyr, one can assume that YMCs must have gone through at least through one crossing time of 1 Gyr. These constraints provide a parameter space for the expected physical properties for YMC progenitors: a mass of  $M_{\text{clump}} > 3 \times 10^4 M_\odot$ , and a size of  $r > r_\Omega = 2GM_{\text{clump}}c_s^{-2}$  for  $M_{\text{clump}} < 8.4 \times 10^4 M_\odot$ , and  $r > r_{\text{vir}} \sim (GM_{\text{clump}}t_{\text{cross}}^2)^{1/3}$  with  $t_{\text{cross}} \sim 1$  Gyr, for  $M_{\text{clump}} > 8.4 \times 10^4 M_\odot$  (Bressert et al. 2012).

YMC progenitors seem to be rare in the Galaxy. Ginsburg et al. (2012) used BOLOCAM dust continuum data to search for YMC progenitors in the Galactic plane. They found three clumps in the solar neighbourhood with characteristics of YMCs progenitors. However, all these clumps showed evidence for active star formation within them. The only YMC progenitor candidate in an early stage of evolution that has been found so far is G0.253+0.016, or ‘The Brick’. ‘The Brick’ (Rathborne et al. 2014a, 2015) is an extremely dense ( $> 10^4 \text{ cm}^{-3}$ ), massive ( $\sim 10^5 M_\odot$ ) and cold ( $\sim 20$  K) molecular cloud (Lis & Carlstrom 1994; Lis & Menten 1998; Lis et al. 2001) located in the harsh environment of the Galactic Centre. The environmental conditions in the Galactic Centre are very different from those of the solar neighbourhood. The column density, gas temperature, velocity dispersion, interstellar radiation field, pressure and cosmic ray ionization rate range from being factors of a few to several orders of magnitude higher compared to the solar neighbourhood (Walmsley et al. 1986; Morris & Serabyn 1996; Ao et al. 2013). The difficulty of finding YMCs progenitors outside the Galactic Centre might also be due to their fast evolution. Indeed Svoboda et al. (2016) estimated the lifetime of YMC progenitors to be 0.03 Myr. Thus, to determine whether starless YMC progenitors exist outside the Galactic Centre we need to extend this search to large samples of molecular clumps in other parts of the Galaxy.

In this paper we present the global properties of 1244 clumps observed in the MALT90 survey. Their wide range of evolutionary stages allows us to analyse their physical properties as a function of their evolution. Section 2 describes the data and the physical properties of the clumps: volume densities, masses, virial masses and bolometric luminosities. In Section 3 we discuss the differences in the physical properties and virial states of the clumps as a function of their dust temperature, used to indicate their evolutionary stage. We identify clumps that could evolve into clusters harbouring high-mass stars, open clusters and YMCs, and we determine the properties of cluster-forming clumps in their earliest evolutionary stage. Finally, Section 4 summarizes the main results of this paper.

## 2 CLUSTER-FORMING CLUMPS WITHIN THE GALAXY

Multiwavelength Galactic plane surveys are extremely valuable for identifying and characterizing, in a systematic way, the properties of cluster-forming clumps within the Galaxy. Infrared (IR) surveys

such as the *Spitzer* 3–8  $\mu\text{m}$  Galactic Legacy Infrared Midplane Extraordinaire (GLIMPSE; Benjamin et al. 2003) and 24  $\mu\text{m}$  MIPS-GAL (Carey et al. 2005) surveys can be used to pinpoint regions of active star formation. In particular, the 4.5  $\mu\text{m}$  *Spitzer* band is commonly used as a tracer of shocked gas, as the enhancement of the emission at this wavelength is associated with the excitation of shocked  $\text{H}_2$  ( $v = 0\text{--}0$ ) S(9,10,11) lines and/or CO ( $v = 1\text{--}0$ ) bandheads. Emission from embedded protostars is usually detected by their emission at 8 and 24  $\mu\text{m}$  due primarily to thermal dust emission (Churchwell et al. 2009). Once the young stars start to ionize their environment, they will form a compact region of ionized hydrogen (H II region) that can be also detected at 8  $\mu\text{m}$  due primarily to bright polycyclic aromatic hydrocarbon (PAH) emission and at 24  $\mu\text{m}$  due to their thermal dust emission.

Galactic plane surveys of the dust continuum emission (e.g. ATLASGAL, Schuller et al. 2009; BGPS, Nordhaus et al. 2008; Hi-Gal, Molinari et al. 2010) provide a complete census of the dust emission towards cluster-forming clumps across the Galaxy (Schuller et al. 2009). Follow-up surveys of their dense molecular gas (e.g. MALT90; Foster et al. 2011, 2013; Jackson et al. 2013; Rathborne et al. 2016; Stephens et al. 2016) have been also key since these data can be used to characterize the gas chemistry and kinematics as well as to provide information about the clumps’ kinematic distances. Moreover, these data are important for determining whether or not adjacent dust continuum peaks are in fact coherent, physically related structures (e.g. Contreras, Rathborne & Garay 2013b; Contreras et al. 2016).

Combined, these surveys provide valuable data sets and allow us to characterize well the star formation potential of hundreds of cluster-forming clumps located throughout the Galaxy.

### 2.1 Identifying cluster-forming clumps

The ATLASGAL dust continuum survey identified more than 10 000 clumps in the Galactic plane from their emission at 870  $\mu\text{m}$  (Contreras et al. 2013a; Urquhart et al. 2014). These clumps cover a large range of evolutionary stages from cold pre-stellar clumps to H II regions. To determine the physical and chemical properties of these star-forming clumps, MALT90 performed targeted observation of 16 spectral lines towards 3557 of the ATLASGAL clumps (Jackson et al. 2013), located between  $-60^\circ$  and  $+20^\circ$  in longitude. The molecular lines observed by the MALT90 survey cover a wide range of excitation energies and critical densities providing a wealth of information about the chemical, physical and kinematic properties of these star-forming clumps (e.g. Sanhueza et al. 2012; Rathborne et al. 2016).

Because dust continuum surveys detect all the emission along the line of sight, it is possible that dust continuum peaks may be due to physically distinct clumps superposed in projection but located at different distances; the kinematic information provided by MALT90 is key to determining whether this is the case or not. The MALT90 catalogue (Rathborne et al. 2016) provides information regarding multiple dense gas molecular clouds along the same line of sight towards the dust continuum emission, differentiating those with single velocity components (S), or two velocity components (A and B). In this work we only used those clumps classified as single (S) that are clearly associated with a single velocity component from the dense gas emission along its line of sight.

MALT90 data were also used to determine the kinematic distances towards the clumps (Whitaker et al., submitted). The kinematic distances were determined based on the flux-weighted average velocity of the emission ( $v_{\text{LSR}}$ ) seen in the HCN(1–0),

HNC(1–0), N<sub>2</sub>H<sup>+</sup>(1–0) and HCO<sup>+</sup>(1–0) transitions. This average velocity was denominated as *consensus velocity* of the clump (see Rathborne et al. 2016). Near and far kinematic distance disambiguation was determined using the H I absorption/emission (e.g. Kolpak et al. 2003; Green & McClure-Griffiths 2011; Wienen et al. 2015). No distances were obtained for clumps within the galactic longitude range  $350^\circ < l < 15^\circ$ . In this region, geometric effects and the large velocity dispersion observed in molecular clouds near the Galactic Centre lead to greater uncertainty in the distance determination. Also, for each clump a reliability factor for its kinematic distance was estimated. This factor has values ranging from –1 (poor) to 10 (most reliable) (see Whitaker et al., submitted). In this work, we only included clumps for which the kinematic distance reliability factor is  $\geq 4$ , to avoid including clumps with kinematic distances not so well determined.

Using the kinematic information provided by MALT90 we determined if two adjacent (in galactic longitude –  $l$  – and latitude –  $b$ ) dust continuum peaks actually arise from the same (more extended) physically coherent structure. We assumed that any column density peak located within  $<1$  arcmin (equivalent to 0.9 pc at 3 kpc) from each other, with a difference in their *consensus velocity*  $<2$  km s<sup>–1</sup> are part of the same physical clump. The values for distance and velocity separation were chosen to make sure that we are not selecting much extended low density emission. Using these parameters, we have identified 44 groups that are connected in the ( $l$ ,  $b$ ,  $v$ ) space.

For these groups of clumps, their areas and mean column densities are derived from the sum and average, respectively, of the clumps that comprise the groups. Table 1 summarizes the physical properties derived for the clumps. In this table the groups of clumps are listed using the MALT90 name of the clump with the highest column density. The column ‘Notes’ in this table shows the individual column density peaks that compose the group of clumps, indicated by their MALT90 name.

## 2.2 Dust temperatures

The dust temperatures of the clumps were derived from a simultaneous fit of the dust temperature and column density to the spectral energy distribution (SED) of the dust continuum emission detected by *Herschel* at 160, 250, 350 and 500  $\mu$ m and by ATLASGAL at 870  $\mu$ m (Guzmán et al. 2015). This fit is done independently for each line of sight, taking into account differences in temperature across the clump better than previous studies (e.g. Faúndez et al. 2004; Hill et al. 2005), in which a single temperature is assumed for the entire clump. However, for clumps with active star formation, it is possible that the column density is underestimated due to temperature gradients along the line of sight. This bias might lead to an underestimation of the column density towards lines of sight associated with large temperature gradients.

In most cases, the consistency between the *Herschel* and ATLASGAL flux measurements allowed for a good fit to dust temperature and column density. However, in some cases ( $<10$  per cent of the sources), in order to perform a good fit, it was necessary to remove some of the ATLASGAL data. In these clumps, the area affected (typically associated with low column density lines of sight) amounts to  $\sim 20$  per cent of the total source area. The errors of the dust temperature and column density best fit were estimated from the projection of confidence ellipses on the dust temperature–column density plane of solutions. Calibration uncertainties in the *Herschel* fluxes are expected to be of the order of 10 per cent, according to the *Herschel* Spectral and Photometric Imaging Receiver (SPIRE) and Photoconductor Array Camera and Spectrometer (PACS) manuals and recommendations. The effect of larger (or smaller) calibration

uncertainties is discussed in Guzmán et al. (2015). For consistency, we decided to adopt 10 per cent calibration uncertainties in this work.

As pre-stellar clumps evolve into stellar clusters, the energy radiation of the embedded protostars will start to heat its surrounding gas and dust. In this scenario, we expect that cold ( $<15$  K) dense ( $>10^4$  cm<sup>–3</sup>) clumps with no evident IR emission represent the earliest stages of cluster formation with no active star formation yet. As the clumps evolve, the UV radiation from the newly formed stars will start to heat and eventually disrupt its environment, increasing the overall temperature of the gas and dust of the clumps. In the IR, these clumps are identified in *Spitzer* GLIMPSE and MIPS GAL images by their enhanced emission at 4.5  $\mu$ m, or with 8 or 24  $\mu$ m point like or extended emission.

Thus, because we expect an increase in the gas and dust temperature of a clump as it evolves, its dust temperature can be used as a proxy for its evolutionary stage. Indeed, Guzmán et al. (2015) showed that the average dust temperature of the MALT90 clumps increased monotonically, from those that showed no IR emission, to those that showed obvious IR emission in *Spitzer* GLIMPSE/MIPSGAL images. This suggests that the classification of the sources into ‘quiescent’, ‘protostellar’, ‘H II regions’ and ‘PDR’ based on their IR emission done by MALT90 followed well the overall evolution of the clumps (see Jackson et al. 2013; Guzmán et al. 2015; Rathborne et al. 2016, for details of this classification). Clumps with no IR evidence of current star formation have dust temperatures of  $\sim 17 \pm 4$  K, while those with IR signatures of active star formation typically have higher dust temperatures of  $\sim 20 \pm 4$  K. Clumps in later evolutionary stages, like those associated with H II regions, typically have dust temperatures of  $>26 \pm 5$  K. Here, the uncertainties represent their internal variance in each group, and not uncertainty of measurements.

Since these results show evidence of a link between the clump dust temperature and the level of star formation activity, in this work we will use the dust temperature of the clump as a systematic measurement of the clump’s evolution. The clumps in our sample have a wide range of dust temperatures, with a median dust temperature of 22.0 K, a minimum of 10 K and a maximum of 44 K (see Table 2).

## 2.3 Clump masses and volume densities

To derive the masses and volume densities of the clumps we used column densities and radii of the clumps derived from the *Herschel* (160, 250, 350 and 500  $\mu$ m) and ATLASGAL (870  $\mu$ m) dust continuum maps and the kinematics distances from Whitaker et al. (submitted). During the fit to the SED, the column density and dust temperature were simultaneously computed, ensuring that the measurements of the column densities are reliable by taking into account the differences of dust temperature within the clump (Guzmán et al. 2015).

The clump effective angular radii given by Guzmán et al. (2015) were obtained after convolving the emission at the higher angular resolutions (18 arcsec at 870  $\mu$ m) by a Gaussian with a size of the beam of lowest angular resolution data (35 arcsec at 500  $\mu$ m). To recover the most accurate measurement of the size of the clump, and to not underestimate the volume density, we computed the beam-corrected effective angular radius of the clump  $\theta_{\text{eff,deconv}}$ . This radius was determined by deconvolving the effective radius of the clump as  $\theta_{\text{eff,deconv}} = \sqrt{\theta_{\text{eff}}^2 - 35^2 \text{ arcsec}^2}$ , where 35 arcsec is the beam size. This size is also comparable to the sizes of the clumps previously determined by Contreras et al. (2013a) and Urquhart et al. (2014) using the ATLASGAL maps. The angular radius was

**Table 1.** Example of the summary table of physical properties derived for the MALT90 clumps with reliable kinematic distances. The full table will be in electronic form. Column 1 shows the MALT90 name of the clumps. Clumps that are part of coherent groups are marked with the † symbol, and in the 'Notes' column is given the information of the associated MALT90 clump. Column 2 shows the kinematic distances (Whitaker et al., submitted). Column 3 corresponds to the consensus velocity. Column 4 corresponds to the mean temperature of the clumps. Column 5 shows the area in parsec. Column 6 shows the effective deconvoluted radius. Column 7 shows the mass derived from the dust continuum emission. Column 8 shows the virial mass. Column 9 shows the virial parameter. Column 10 corresponds to the volume density. Column 11 shows the bolometric luminosity derived from the extrapolation of the dust continuum emission, and Column 12 shows the clumps associated in  $l$ - $b$ - $v$  space.

Name	Dist. (kpc)	$\Delta v$ (km s <sup>-1</sup> )	Temp. (K)	Area (pc <sup>2</sup> )	$R_{\text{eff}}$ (pc)	Mass (10 <sup>3</sup> M <sub>⊙</sub> )	Virial mass (10 <sup>3</sup> M <sub>⊙</sub> )	$\alpha$	$n(\text{H}_2)$ (10 <sup>4</sup> cm <sup>-3</sup> )	$L_{\text{bol}}$ (10 <sup>2</sup> L <sub>⊙</sub> )	Notes
AGAL305.226+00.274_S†	4.9	1.7 ± 0.0	20 ± 2	6.4 ± 1.7	1.4 ± 0.2	7.2 ± 2.0	4.6 ± 0.6	0.6 ± 0.2	1.5 ± 0.2	390.8 ± 117.0	AGAL305.237+00.286_S
AGAL305.234+00.262_S	5.3	1.5 ± 0.0	23 ± 2	0.0 ± 0.0	0.1 ± 0.0	2.5 ± 0.9	0.2 ± 0.0	0.1 ± 0.0	1133.4 ± 205.7	4.1 ± 1.7	
AGAL305.236+00.022_S	7.1	1.6 ± 0.0	18 ± 2	14.1 ± 3.5	2.1 ± 0.3	10.3 ± 2.6	4.2 ± 0.6	0.4 ± 0.1	0.4 ± 0.1	428.4 ± 136.1	
AGAL305.242+00.041_S	6.8	1.6 ± 0.1	16 ± 2	0.8 ± 0.3	0.5 ± 0.1	3.9 ± 1.2	1.1 ± 0.2	0.3 ± 0.1	11.6 ± 1.9	10.8 ± 4.0	
AGAL305.242+00.276_S	4.9	1.4 ± 0.0	22 ± 6	1.5 ± 0.6	0.7 ± 0.1	0.8 ± 0.3	1.1 ± 0.2	1.3 ± 0.5	1.0 ± 0.2	309.2 ± 131.0	
AGAL305.257+00.004_S	6.8	1.4 ± 0.2	37 ± 2	1.5 ± 0.5	0.7 ± 0.1	0.2 ± 0.1	1.2 ± 0.3	7.0 ± 3.0	0.2 ± 0.0	286.2 ± 105.6	
AGAL305.259+00.326_S	5.2	1.6 ± 0.0	18 ± 2	0.2 ± 0.1	0.2 ± 0.0	1.3 ± 0.5	0.5 ± 0.1	0.4 ± 0.2	37.9 ± 6.9	2.3 ± 1.0	
AGAL305.271+00.029_S	6.7	2.0 ± 0.1	22 ± 1	2.4 ± 1.0	0.9 ± 0.2	1.0 ± 0.4	2.9 ± 0.6	2.8 ± 1.3	0.6 ± 0.1	65.2 ± 28.7	
AGAL305.271+00.009_S	6.7	1.9 ± 0.0	27 ± 2	12.0 ± 4.7	2.0 ± 0.4	9.3 ± 3.6	5.7 ± 1.1	0.6 ± 0.3	0.5 ± 0.1	1687.7 ± 738.1	
AGAL305.272+00.296_S	5.6	1.2 ± 0.1	26 ± 1	3.6 ± 1.3	1.1 ± 0.2	2.7 ± 1.0	1.2 ± 0.2	0.4 ± 0.2	0.9 ± 0.2	258.9 ± 104.6	
AGAL305.346+00.212_S	4.9	2.7 ± 0.1	32 ± 1	2.2 ± 0.8	0.8 ± 0.1	2.6 ± 0.9	4.9 ± 0.9	1.8 ± 0.7	1.8 ± 0.3	959.3 ± 384.7	
AGAL305.357+00.202_S	5.8	2.3 ± 0.1	32 ± 1	3.1 ± 1.1	1.0 ± 0.2	3.6 ± 1.3	4.4 ± 0.8	1.2 ± 0.5	1.5 ± 0.3	1315.4 ± 531.6	
AGAL305.361+00.186_S	5.5	1.4 ± 0.0	26 ± 1	4.3 ± 1.5	1.2 ± 0.2	5.5 ± 1.9	1.8 ± 0.3	0.3 ± 0.1	1.4 ± 0.3	1887.1 ± 758.1	
AGAL305.367+00.212_S	6.5	1.8 ± 0.1	28 ± 1	5.9 ± 2.2	1.4 ± 0.3	5.6 ± 2.1	3.6 ± 0.8	0.6 ± 0.3	0.9 ± 0.2	1402.3 ± 598.2	
AGAL305.381+00.241_S	5.6	1.9 ± 0.3	23 ± 4	1.5 ± 0.5	0.7 ± 0.1	1.1 ± 0.4	2.0 ± 0.8	1.8 ± 0.9	1.4 ± 0.3	71.3 ± 28.7	
AGAL305.509+00.367_S	5.9	1.7 ± 0.3	23 ± 2	1.7 ± 0.6	0.7 ± 0.1	0.4 ± 0.1	1.7 ± 0.8	4.1 ± 2.3	0.4 ± 0.1	51.7 ± 21.0	
AGAL305.512+00.334_S	4.9	1.1 ± 0.2	21 ± 1	1.0 ± 0.4	0.6 ± 0.1	0.2 ± 0.1	0.5 ± 0.2	2.4 ± 1.3	0.5 ± 0.1	18.4 ± 7.7	
AGAL305.512+00.349_S	5.6	1.4 ± 0.2	24 ± 2	2.0 ± 0.7	0.8 ± 0.1	0.5 ± 0.2	1.2 ± 0.4	2.5 ± 1.3	0.4 ± 0.1	79.0 ± 31.7	
AGAL305.521+00.361_S	6.0	2.4 ± 0.3	27 ± 3	1.7 ± 0.6	0.7 ± 0.1	0.3 ± 0.1	3.4 ± 1.1	11.6 ± 5.7	0.3 ± 0.1	87.0 ± 35.5	
AGAL305.536+00.019_S	5.4	2.4 ± 0.3	33 ± 3	1.1 ± 0.4	0.6 ± 0.1	0.2 ± 0.1	2.7 ± 0.8	13.8 ± 6.3	0.4 ± 0.1	151.1 ± 60.6	
AGAL305.539+00.339_S	6.5	1.3 ± 0.1	24 ± 2	5.0 ± 1.9	1.3 ± 0.2	2.6 ± 1.0	1.7 ± 0.4	0.7 ± 0.3	0.5 ± 0.1	279.2 ± 119.2	
AGAL305.549+00.004_S	5.9	1.7 ± 0.2	33 ± 2	1.3 ± 0.5	0.7 ± 0.1	0.3 ± 0.1	1.6 ± 0.4	5.0 ± 2.3	0.5 ± 0.1	248.6 ± 101.0	
AGAL305.554+00.011_S	3.8	1.6 ± 0.3	26 ± 1	2.1 ± 1.2	0.8 ± 0.2	0.8 ± 0.4	1.7 ± 0.7	2.2 ± 1.6	0.6 ± 0.2	260.4 ± 158.2	
AGAL305.562+00.014_S	5.4	1.2 ± 0.1	32 ± 1	3.8 ± 1.3	1.1 ± 0.2	1.9 ± 0.7	1.3 ± 0.3	0.7 ± 0.3	0.6 ± 0.1	852.5 ± 342.7	
AGAL305.574+00.342_S	4.9	1.5 ± 0.3	21 ± 1	0.5 ± 0.1	0.4 ± 0.0	0.1 ± 0.0	0.8 ± 0.3	10.3 ± 4.2	0.4 ± 0.0	5.0 ± 1.3	
AGAL305.581+00.037_S	6.0	2.0 ± 0.3	34 ± 1	0.7 ± 0.3	0.5 ± 0.1	0.4 ± 0.1	1.6 ± 0.5	4.4 ± 2.1	1.4 ± 0.3	54.4 ± 22.1	
AGAL305.589+00.462_S	3.8	1.3 ± 0.2	16 ± 2	0.1 ± 0.1	0.2 ± 0.1	0.4 ± 0.2	0.3 ± 0.1	0.6 ± 0.4	22.7 ± 6.6	0.9 ± 0.5	
AGAL305.777+00.249_S	2.7	1.0 ± 0.1	19 ± 3	0.1 ± 0.1	0.2 ± 0.1	0.2 ± 0.1	0.2 ± 0.1	0.8 ± 0.6	14.9 ± 4.8	1.4 ± 0.9	
AGAL305.794+00.096_S	5.0	1.4 ± 0.1	16 ± 2	1.3 ± 0.5	0.6 ± 0.1	1.6 ± 0.5	1.1 ± 0.2	0.7 ± 0.3	2.4 ± 0.4	9.6 ± 3.9	

**Table 2.** Summary of physical properties derived for the full sample and divided by dust temperature of the clumps.

Sample	Per cent	Temperature (K)				$R_{\text{eff}}$ (pc)			
		Median	$\sigma_{\text{Med.}}$	Min	Max	Median	$\sigma_{\text{Med.}}$	Min	Max
Full sample		22	5.9	10	44	0.82	0.6	0.03	8.7
$T_d \leq 15$	12	14	1.5	10	15	0.64	0.46	0.071	3.1
$15 < T_d \leq 20$	30	18	1.6	15	20	0.74	0.55	0.033	5.7
$20 < T_d \leq 25$	30	22	1.9	20	25	0.83	0.59	0.03	8.7
$25 < T_d \leq 30$	19	27	1.5	25	30	0.93	0.66	0.061	4.4
$30 < T_d$	9	33	2.8	30	44	0.77	0.54	0.041	3.9
HMPC	78	21	5.9	10	44	0.92	0.68	0.03	8.7
YMC	0.4	20	3	14	23	3.6	1.6	2.5	5.7
OC	67	22	5.9	10	44	0.97	0.58	0.4	3.7
Cold, dark 70 $\mu\text{m}$	2	14	2.3	10	19	0.7	0.37	0.061	2.1
Sample		Area ( $\text{pc}^2$ )				Mass ( $M_\odot$ )			
		Median	$\sigma_{\text{Med.}}$	Min	Max	Median	$\sigma_{\text{Med.}}$	Min	Max
Full sample		2.1	2.5	0.0029	$2.4e+02$	$1e+03$	$1.2e+03$	0.99	$1e+05$
$T_d \leq 15$		1.3	1.5	0.016	31	$9.8e+02$	$1.2e+03$	17	$4.1e+04$
$15 < T_d \leq 20$		1.7	2.2	0.0033	$1e+02$	$1e+03$	$1.1e+03$	12	$9.9e+04$
$20 < T_d \leq 25$		2.2	2.5	0.0029	$2.4e+02$	$1.1e+03$	$1.3e+03$	2.3	$1e+05$
$25 < T_d \leq 30$		2.7	3.2	0.012	61	$1.2e+03$	$1.4e+03$	18	$4.5e+04$
$30 < T_d$		1.9	2.1	0.0052	47	$6.3e+02$	$8e+02$	0.99	$1.6e+04$
HMPC		2.7	3.3	0.0029	$2.4e+02$	$1.5e+03$	$1.6e+03$	12	$1e+05$
YMC		41	31	20	$1e+02$	$6.7e+04$	$4.8e+04$	$3.3e+04$	$1e+05$
OC		3	3.1	0.5	42	$1e+03$	$1.1e+03$	37	$1.4e+04$
Cold, dark 70 $\mu\text{m}$		1.5	1.3	0.012	14	$1.2e+03$	$1.2e+03$	$1.9e+02$	$2.5e+04$
Sample		Virial mass ( $M_\odot$ )				$\alpha$			
		Median	$\sigma_{\text{Med.}}$	Min	Max	Median	$\sigma_{\text{Med.}}$	Min	Max
Full sample		$1.2e+03$	$1.1e+03$	21	$3.8e+04$	1.1	0.95	0.036	$2.4e+02$
$T_d \leq 15$		$8.4e+02$	$8.5e+02$	75	$1.3e+04$	0.8	0.71	0.036	30
$15 < T_d \leq 20$		$9.9e+02$	$9.4e+02$	21	$3.7e+04$	0.94	0.86	0.053	20
$20 < T_d \leq 25$		$1.2e+03$	$9.6e+02$	21	$3.4e+04$	0.96	0.87	0.071	$1.4e+02$
$25 < T_d \leq 30$		$1.5e+03$	$1.5e+03$	63	$3.8e+04$	1.3	1.1	0.11	32
$30 < T_d$		$1.6e+03$	$1.6e+03$	75	$1.9e+04$	2.1	2	0.19	$2.4e+02$
HMPC		$1.3e+03$	$1.2e+03$	21	$3.8e+04$	0.82	0.64	0.036	19
YMC		$1.1e+04$	$3.3e+03$	$6.6e+03$	$2.5e+04$	0.21	0.067	0.067	0.36
OC		$1.3e+03$	$1.1e+03$	21	$3.8e+04$	1.3	0.95	0.11	41
Cold, dark 70 $\mu\text{m}$		$1.1e+03$	$1.1e+03$	51	$3.9e+03$	0.51	0.43	0.15	7.5
Sample		$n(\text{H}_2)$ ( $\text{g cm}^{-3}$ )				$L_{\text{bol}}$ ( $L_\odot$ )			
		Median	$\sigma_{\text{Med.}}$	Min	Max	Median	$\sigma_{\text{Med.}}$	Min	Max
Full sample		$6.8e+03$	$6.7e+03$	$4.9e+02$	$1.1e+07$	$4.9e+03$	$6.8e+03$	4.9	$1.4e+06$
$T_d \leq 15$		$1.2e+04$	$1.1e+04$	$1.6e+03$	$1.2e+06$	$4.3e+02$	$5.3e+02$	6.3	$3.1e+04$
$15 < T_d \leq 20$		$7.3e+03$	$7e+03$	$8.7e+02$	$6e+06$	$1.7e+03$	$2.1e+03$	5.5	$5e+05$
$20 < T_d \leq 25$		$6.2e+03$	$6.1e+03$	$4.9e+02$	$1.1e+07$	$7.3e+03$	$8.8e+03$	4.9	$1.1e+06$
$25 < T_d \leq 30$		$5.2e+03$	$5.6e+03$	$6.6e+02$	$2.5e+06$	$1.9e+04$	$2.2e+04$	67	$6.9e+05$
$30 < T_d$		$4.5e+03$	$4.3e+03$	$7.5e+02$	$6.1e+05$	$3e+04$	$3.8e+04$	62	$1.4e+06$
HMPC		$8.5e+03$	$8.6e+03$	$4.9e+02$	$1.1e+07$	$6.3e+03$	$8.7e+03$	4.9	$1.4e+06$
YMC		$5.9e+03$	$3.6e+03$	$2.2e+03$	$8.6e+03$	$1.8e+05$	$2.2e+05$	$3.1e+04$	$5.4e+05$
OC		$4.7e+03$	$3.9e+03$	$6.6e+02$	$1.1e+05$	$6.9e+03$	$8.8e+03$	44	$1e+06$
Cold, dark 70 $\mu\text{m}$		$1e+04$	$5.8e+03$	$2.5e+03$	$6e+06$	$8.5e+02$	$9.1e+02$	5.5	$6.9e+03$



converted then into physical radius using the distance to the sources via  $R_{\text{eff}} = \theta_{\text{eff,deconv}} D / 206\,265$ , with  $D$  the distance to the clump in pc and  $\theta_{\text{eff,deconv}}$  in arcsec.

The total mass for each clump was derived using the expression

$$M = 4.8 \times 10^3 \left( \frac{N(\text{H}_2)_{\text{mean}}}{\text{g cm}^{-2}} \right) \left( \frac{A}{\text{pc}^2} \right) \mu_{\text{H}_2} M_{\odot}, \quad (1)$$

where  $N(\text{H}_2)_{\text{mean}}$  is the mean column density within the clump derived from the dust continuum emission assuming a gas-to-dust mass ratio of 100,  $A$  is its area, with  $A = \pi R_{\text{eff}}^2$  and  $\mu_{\text{H}_2} = 2.8$  is the mean molecular weight. The derived clump masses range from a few  $M_{\odot}$  to  $\sim 10^5 M_{\odot}$ , with a median value of  $1.0 \times 10^3 M_{\odot}$  (see Table 2).

The volume density,  $n(\text{H}_2)$ , was computed from the total mass and the volume ( $V = \frac{4}{3}\pi R_{\text{eff}}^3$ ), assuming a spherical shape and a molecular weight of  $\mu_{\text{H}_2} = 2.8$  (see Table 2). The volume density of clumps ranges from  $4.9 \times 10^2$  to  $1.1 \times 10^7 \text{ cm}^{-3}$ . The median volume density is  $\sim 6.8 \times 10^3 \text{ cm}^{-3}$ .

### 2.3.1 Virial mass and gravitational state of the clumps

The gravitational state of each clump was determined by measuring its virial parameter,  $\alpha$ , defined as the ratio of the clump's virial mass to its total mass. The parameter  $\alpha$  is commonly used to indicate whether or not a clump is gravitationally bound. If  $\alpha > 1$ , then the clump contains more kinetic energy than gravitational energy, and is likely to expand unless confined by an external mechanism. In this case, the clump is considered to be unbound. A value of  $\alpha < 1$  suggests that the kinetic energy is insufficient to support the clump against its gravitational energy, and the clump will likely undergo collapse. In this case, in the absence of magnetic fields, the clump is considered to be bound.

The virial mass for each clump was calculated via (Bertoldi & McKee 1992)

$$M_{\text{vir}} = \frac{5}{G a_1 a_2} R_{\text{eff}} \sigma^2, \quad (2)$$

where  $G$  is the gravitational constant;  $a_1 = (1 - p/3)/(1 - 2p/5)$ , for  $p < 2.5$ , which is a correction term that depends on the density gradient with power-law index  $p$  (Dunham et al. 2010), which we assume to be  $p = 1.8$  (Mueller et al. 2002); and  $a_2$  is a correction term that depends on the clump shape. For simplicity we assume the clumps are spherical, thus  $a_2$  is 1 for all clumps.  $R_{\text{eff}}$  is the effective radius of the clumps and  $\sigma$  is the clump's velocity dispersion.

The velocity dispersion for each clump was determined from the  $\text{N}_2\text{H}^+$  or HNC line-widths obtained from the MALT90 molecular line emission (Rathborne et al. 2016), as  $\sigma = \Delta v / 2.35$ , assuming that the lines have Gaussian profiles. We used these molecules as they are good density tracers for the clumps, and the line-widths are not so severely affected by the dynamics of the source (e.g. outflows, infall profiles). If the  $\text{N}_2\text{H}^+$  emission was detected towards the clump and the  $\text{N}_2\text{H}^+$  hyperfine components were fitted, the  $\text{N}_2\text{H}^+$  line-widths were used to determine the velocity dispersion of the clump. However, if the fit to the  $\text{N}_2\text{H}^+$  hyperfine components was poor, we used instead the HNC line-widths to determine the velocity dispersion of the clump. For 91 clumps, there was no reliable detection of either  $\text{N}_2\text{H}^+$  or HNC, thus, we did not compute the virial mass for these clumps and these clumps are not considered in the rest of the analysis. Thus, the total number of clumps included in our sample is 1244.

The derived values for each clump's velocity dispersion, virial mass and virial parameter are summarized in Table 2. The velocity

dispersion ranges from 0.2 to 5  $\text{km s}^{-1}$ . We find typical values for the virial mass ranging from 21 to  $3.8 \times 10^4 M_{\odot}$ , with a median value of  $1.2 \times 10^3 M_{\odot}$ . The virial parameter ranges from 0.04 to 240 with a median of 1.1.

### 2.3.2 Bolometric luminosity

The bolometric luminosity of a clump provides information about the star formation activity within the clump. If the clump is in an early evolutionary stage with no embedded protostars, then the clump will have a low bolometric luminosity. On the other hand, if star formation has already started within the clump, then we expect that the radiation of the newly formed stars will increase its overall bolometric luminosity. Thus, the bolometric luminosity can also be used to evaluate the level of active star formation within a clump.

We calculate the bolometric luminosity ( $L_{\text{bol}}$ ) extrapolated from the dust emission detected towards the clumps. Here we assume that all of the radiation coming from the clump can be described by a grey body with the opacity law given by  $\kappa_v = 0.14 (250 \mu\text{m} / \lambda_{\mu\text{m}})^{\beta} \text{ cm}^2 \text{ g}^{-1}$ , and any internal optical or ultraviolet (UV) radiation is reprocessed by dust and emerges in the far-IR/submm. The extrapolation of the fit to the dust thermal emission is expressed in terms of the column density, solid angle and dust temperature of the clumps via

$$L_{\text{bol}} = 40 L_{\odot} \left( \frac{T}{10 \text{ K}} \right)^{5.7} \left( \frac{D}{\text{kpc}} \right)^2 \left( \frac{\Omega_s}{\text{arcmin}^2} \right) \times \left( \frac{N_{\text{gas}}}{\text{g cm}^{-2}} \right) C(T, N_{\text{gas}}), \quad (3)$$

where  $D$  is the kinematic distance;  $\Omega_s$  is the solid angle;  $N_{\text{gas}} = N_{\text{H}_2} \mu_{\text{gas}} m_{\text{H}}$  is the mean gas column density, with  $\mu_{\text{H}_2} = 2.8$  the mean molecular weight for the molecular gas and  $m_{\text{H}}$  the hydrogen mass;  $T$  is the dust temperature and  $C(T, N_{\text{H}_2})$  is a correction factor that takes into account opacity and dust temperature effects given by

$$C(T, N_{\text{gas}}) = (1 + 1.6 \times 10^{-3} N_{\text{gas}} T^{1.7})^{-1}. \quad (4)$$

Appendix A provides a more detailed explanation on the derivation of  $L_{\text{bol}}$ .

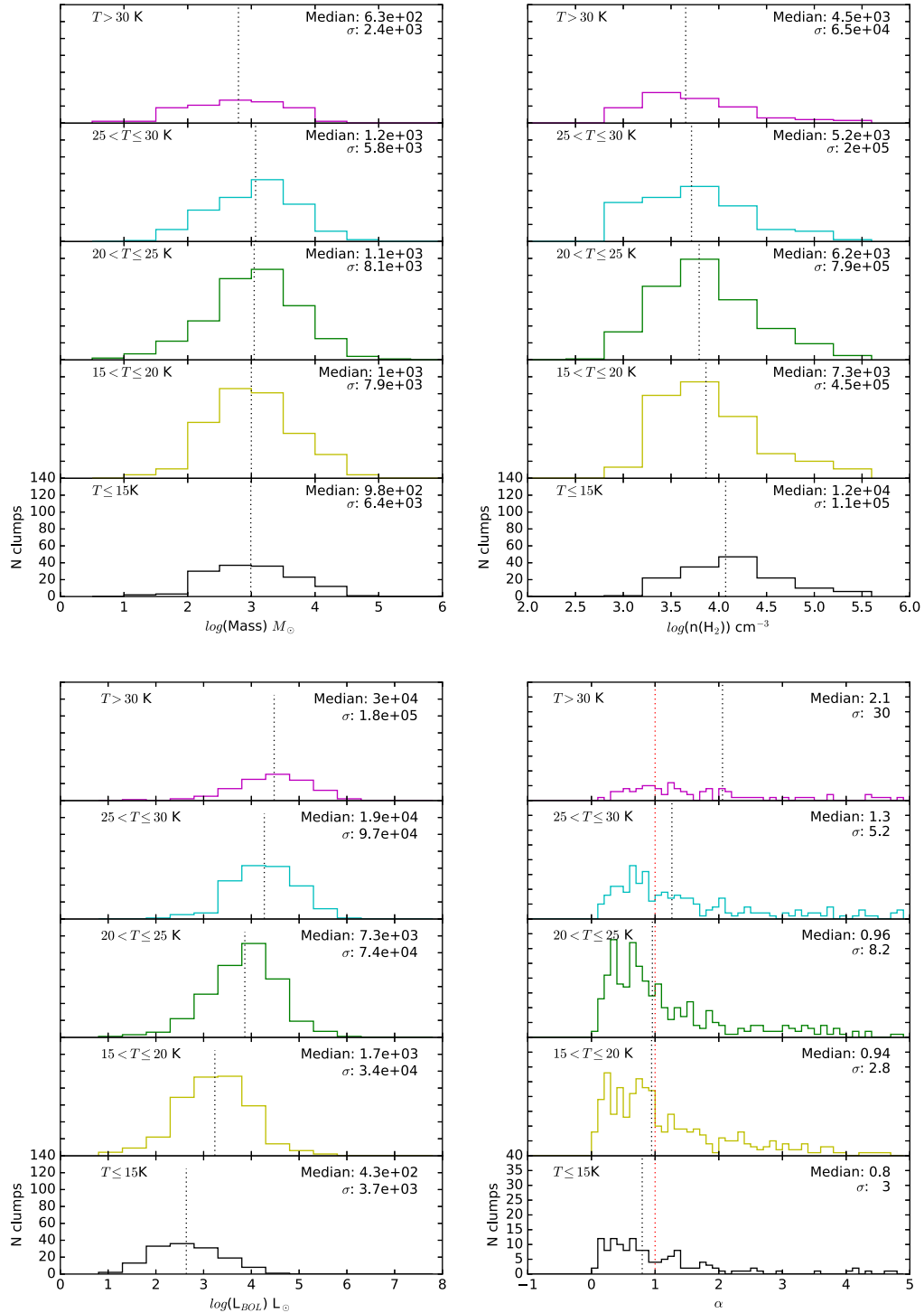
The bolometric luminosity of the clumps ranges from 4.9 to  $1.4 \times 10^6 L_{\odot}$ , with a median value of  $4.9 \times 10^3 L_{\odot}$  (see Table 2).

## 3 DISCUSSION

### 3.1 Evolution of the physical properties

We use the dust temperature of the clumps to trace the evolution in their physical parameters (mass, volume density, bolometric luminosity and virial state). To do this, we separated the clumps into dust temperature bins of 5 K, between 15 and 30 K. Clumps with dust temperature  $\leq 20$  K represent clumps in their earliest evolutionary stages, clumps with dust temperatures in the range  $20 < T_{\text{dust}} \leq 25$  K are those where star formation has already started, and clumps with dust temperatures  $> 25$  K are those at the most evolved evolutionary stages, usually associated with H II regions and PDRs. Table 2 shows the range of values for the full sample and for each of the dust temperature bins, and the percentage of clumps within each dust temperature bin.

Fig. 1 shows the mass, volume density, bolometric luminosity and virial state distribution of the clumps separated by their dust temperatures. The median mass for the clumps ranges from  $\sim 0.6 \times 10^3$



**Figure 1.** Histograms of clump mass, volume density, bolometric luminosity and virial parameter for the indicated dust temperature bins. In each plot the black dotted line shows the median value of the histogrammed quantity. This figure shows that the median value of the mass does not change significantly for clumps with different dust temperatures, having an median mass of  $\sim 10^3 \pm 10^2 M_{\odot}$  (upper left-hand panel). The upper right-hand panel shows the volume density of the clumps for the same dust temperature bins. This figure shows that the volume density decreases for clumps with higher dust temperature, which is consistent with the evolution of the clumps in which clumps evolve thus diminishing their average volume density. The lower left-hand panel shows the bolometric luminosity distributions for the different dust temperature bins. We see an increase of the bolometric luminosity with the evolution of the clumps, also consistent with the evolution of the clumps. The lower right-hand panel shows the distribution of the virial parameter  $\alpha$  for the different dust temperature bins. In this plot the red dotted line shows  $\alpha = 1$ . The virial parameter median value is  $\lesssim 1$  for clumps with dust temperature  $< 25$  K and increases to  $\alpha > 1$  for clumps with higher dust temperature. This is also consistent with the evolution of the clumps, where clumps are gravitationally bound at the onset of star formation and become unbound when they evolve.

**Table 3.** KS statistical test values for the distribution of mass, density, bolometric luminosity and virial parameter as function of dust temperature bins.

Sample	Mass		$n(\text{H}_2)$		$L_{\text{bol}}$		$\alpha$	
	KS value	$p$ value	KS value	$p$ value	KS value	$p$ value	KS value	$p$ value
$T_d \leq 15$	4.41e-02	9.60e-01	2.49e-01	1.49e-07	4.87e-01	1.41e-27	1.40e-01	1.15e-02
$15 < T_d \leq 20$	3.67e-02	8.28e-01	5.63e-02	3.15e-01	2.49e-01	4.40e-16	5.78e-02	2.84e-01
$20 < T_d \leq 25$	3.84e-02	7.80e-01	3.50e-02	8.66e-01	1.30e-01	1.05e-04	5.11e-02	4.29e-01
$25 < T_d \leq 30$	6.63e-02	3.39e-01	1.13e-01	1.18e-02	3.24e-01	8.23e-19	7.98e-02	1.54e-01
$30 < T_d$	1.62e-01	6.83e-03	1.52e-01	1.29e-02	4.02e-01	1.12e-15	2.55e-01	1.46e-06

to  $1.2 \times 10^3 M_\odot$  for all the evolutionary stages, and is similar to the median value of the full sample of  $1.0 \times 10^3 M_\odot$ .

For each dust temperature bin, we computed a Kolmogorov–Smirnov (KS) test of the mass distribution with respect to the full sample of clumps. Table 3 shows the results of the KS value and probability for each physical parameter. We find that for clumps with  $T_{\text{dust}} \leq 30$  K, the mass distribution in each bin is similar to the mass distribution of the full sample. For clumps with  $T_{\text{dust}} > 30$  K the KS test shows that the mass distribution is different from the mass distribution of the full sample, which might be due to the lower number of clumps with these dust temperatures. The median values of the mass and the KS test suggest that the mass distribution of the clumps remains fairly constant regardless of their dust temperature, thus, we are tracing similar mass ranges for all evolutionary stages.

KS tests on the mean volume density distribution shows that the density distribution for each dust temperature bin is different from the distribution of full sample (see Table 3). This suggests that the volume density of a clump changes with its evolution. We found that the median volume density of clumps decreases with the dust temperature. This is consistent with the scenario where at later evolutionary stages, and thus higher dust temperatures, the clumps have already formed protostars and have started to expand, thereby increasing their size and decreasing their mean volume density.

The bolometric luminosity of the clumps also increases with the dust temperature of the clumps. This change is confirmed by the KS test between the distribution of bolometric luminosity for each dust temperature bin and the full sample (see Table 3). We found that the dust bolometric luminosity of the clumps increases with dust temperature as expected, given that the bolometric luminosity is a steeply increasing function of the dust temperature of the clump. This is also consistent with the fact that clumps with higher temperature likely already have embedded protostars, which are heating their environment, increasing the overall bolometric luminosity of the clumps.

We now analyse the virial parameters of the clumps. KS tests show that the distributions for clumps with dust temperature  $\leq 15$  K and clumps with  $T_{\text{dust}} > 25$  K are different from the distribution exhibited by the full sample. However, for clumps with dust temperature in the range  $15 < T_{\text{dust}} \leq 25$  K the distribution of the virial parameter is similar to the distribution of the full sample. Thus, the virial parameter also increases with the evolution of the clumps. For clumps with dust temperatures  $\leq 25$  K, the median values of  $\alpha$  are  $\lesssim 1$ , while for clumps with dust temperatures  $> 25$  K, the median value of  $\alpha$  is greater than 1. This trend is also consistent with the evolution of the clumps. When the clumps have a lower dust temperature, they are more bound. Once star formation has begun, the protostars will start to heat and ionize their environment. The temperature of the clump will increase, and the protostar will start dissipating its natal gas. The feedback from the newly formed stars (e.g. outflows and winds) will increase the turbulence and therefore the virial mass. At this stage the clump as a whole will become more gravitationally unbound.

### 3.2 Free-fall time and dynamical time of the clumps

We determined the degree of gravitational instability of the clumps as a function of the clumps masses. For this we determined the free-fall time,  $t_{\text{ff}} = \sqrt{3\pi/32G\rho}$ , and the dynamical crossing time or dynamical time,  $t_{\text{dyn}} = R_{\text{eff}}/\sigma$  (Tan, Krumholz & McKee 2006; McKee & Ostriker 2007), for the clumps. A clump will become gravitationally unstable if the free-fall time is less than its dynamical time. Moreover, since the ratio  $t_{\text{eff}}/t_{\text{dyn}} \propto \sqrt{\alpha}$ , as  $t_{\text{ff}}/t_{\text{dyn}}$  goes below 1, clumps will become subvirial.

The free-fall time is relatively insensitive to the value of the clump mass, but the dynamical time is sensitive to the value of the mass. Because gravitational collapse will start for those clumps for which the free-fall time is less than the dynamical time, this result suggests that clumps with higher masses are more susceptible to gravitational instabilities and they evolve faster than their lower mass counterparts (see Fig. 2). This is in agreement with the theoretical picture for high-mass star formation, where clumps with higher mass evolve faster, quickly disrupting their natal environment (e.g. Tan et al. 2014).

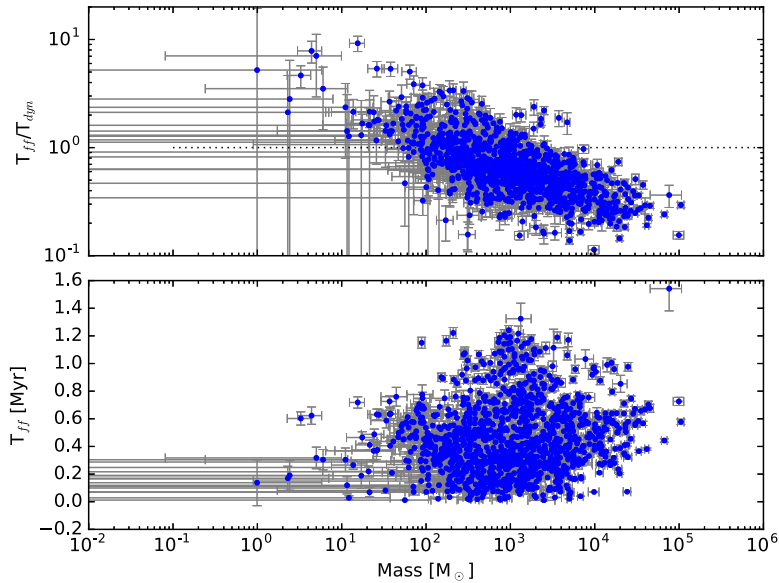
While most of the clumps have  $t_{\text{ff}}/t_{\text{dyn}} < 1$  and thus are likely to collapse, some clumps have  $t_{\text{ff}}$  greater than their  $t_{\text{dyn}}$ . These clumps can be divided into two categories: (1) clumps where star formation has already begun (dust temperature  $> 20$  K), with high virial parameter ( $\langle \alpha \rangle = 5.6 \pm 1.8$ ) due to activity from the stars (e.g. outflows, winds that increase the turbulence within the clump) and thus high  $t_{\text{ff}}/t_{\text{dyn}}$  ( $\langle t_{\text{ff}}/t_{\text{dyn}} \rangle = 1.4 \pm 0.2$ ), and (2) clumps in early stages of evolution (dust temperature  $\leq 20$  K) that have low volume density ( $\langle n(\text{H}_2) \rangle = 6.4 \pm 0.5 \times 10^3 \text{ g cm}^{-3}$ ), low bolometric luminosities ( $\langle L_{\text{bol}} \rangle = 4.7 \pm 0.9 L_\odot$ ) and are unbound ( $\langle \alpha \rangle = 5.0 \pm 1.3$ ); thus, they might not collapse to form stars.

To determine the potential of collapse for the clumps that are in an early evolutionary stage we analysed the ratio of the free-fall time to the dynamical time ( $t_{\text{ff}}/t_{\text{dyn}}$ ) for the clumps with dust temperatures  $\leq 20$  K (See Fig. 3). We determined the fraction of clumps that have  $t_{\text{ff}}/t_{\text{dyn}} < 1$  compared to the total number of clumps for any given volume density. For volume densities  $< 2.5 \times 10^4 \text{ cm}^{-3}$ , 82 per cent of the clumps in our sample have values of  $t_{\text{ff}}/t_{\text{dyn}} < 1$  and thus they are prone to gravitational collapse, while for volume densities  $> 2.5 \times 10^4 \text{ cm}^{-3}$  the number of clumps that have values of  $t_{\text{ff}}/t_{\text{dyn}} < 1$  increases to a 97 per cent. This value of volume density agrees with previous studies of solar neighbourhood clouds, which suggest that star formation only occurs when the volume density exceeds a threshold density of  $\sim 10^4 \text{ cm}^{-3}$  (e.g. Lada et al. 2012).

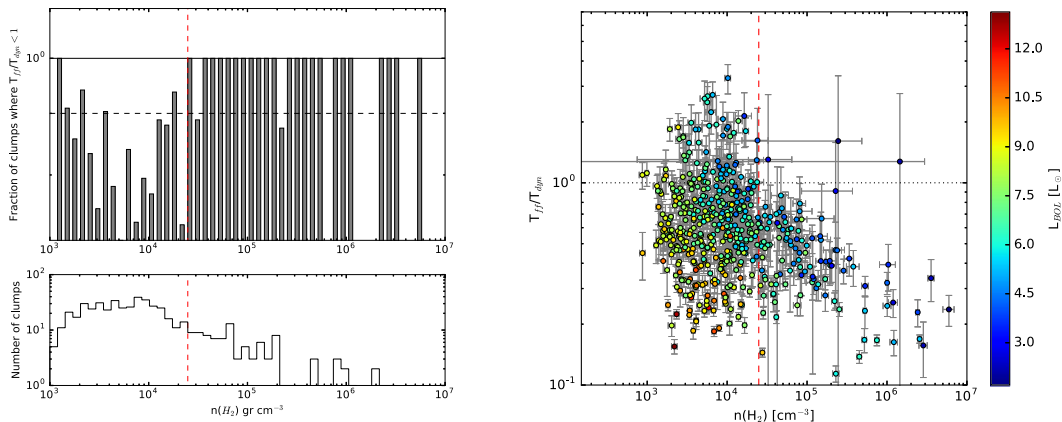
### 3.3 High-mass star and cluster-forming potential of the clumps

To determine which clumps have sufficient mass and density to form stars and stellar clusters, we analysed the mass and radius of the clumps. Fig. 4 shows the mass–radius relationship for the clumps, colour coded by dust temperature bins. By fitting the mass





**Figure 2.** Upper panel: ratio of the free-fall to the dynamical time versus mass of the clumps, the dotted line marks where the free-fall times are equal to the dynamical time. The crosses show the uncertainties in the measurements, which are mostly due to uncertainties in the kinematic distance. In this plot we can see that the clumps with higher mass evolve faster. Lower panel: free-fall time versus mass of the clumps. This plot shows that while there is a large dispersion in the free-fall time, it remains fairly constant with the mass of clump.



**Figure 3.** Left, upper panel: fraction of clumps with  $\tau_{ff}/\tau_{dyn} < 1$  as a fraction of the total number of clumps. The solid black line marks the ratio =1 and the dashed black line marks the ratio equals 0.9. Left, lower panel: number of clumps as a function of volume density. Right-hand panel: ratio of the free-fall time over dynamical time ( $\tau_{ff}/\tau_{dyn}$ ) versus volume density for the clumps with dust temperature  $\leq 20$  K. Colours show the bolometric luminosity of the clumps. Grey crosses show the uncertainty in the measurements. The dotted black line marks  $\tau_{ff} = \tau_{dyn}$ , clumps below this line are more susceptible to instability and collapse. In all the plots the dashed red lines mark the volume density for which the ratio starts to be >90 per cent,  $n(H_2) = 2.5 \times 10^4 \text{ cm}^{-3}$ , above which nearly all the clumps are prone to collapse.

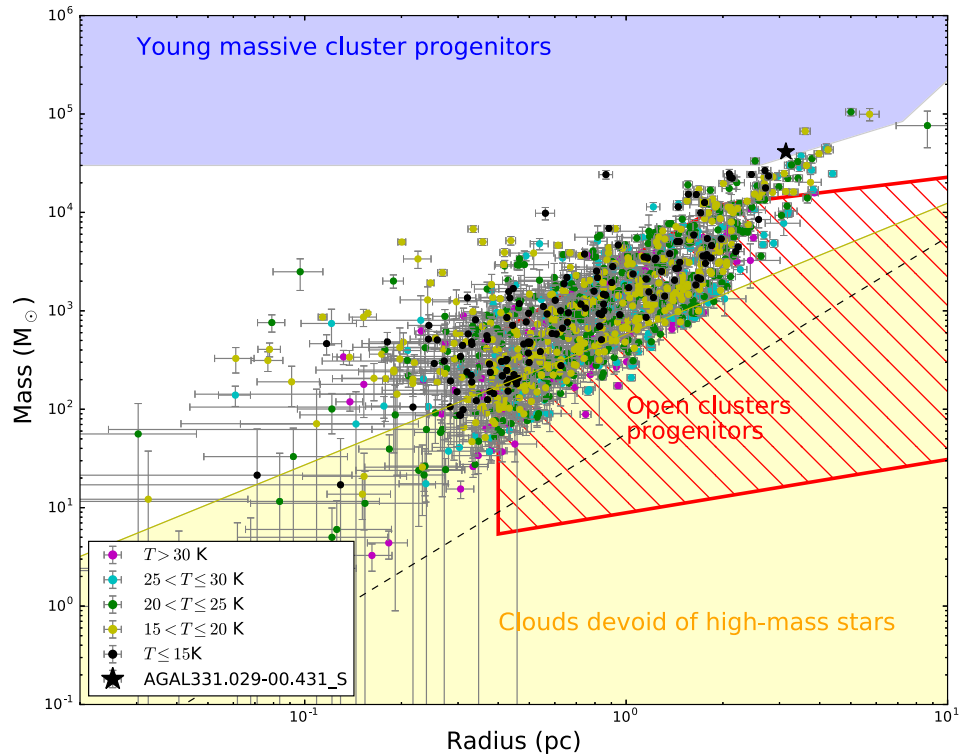
and radius of the clumps, we found that they follow the relationship  $M = 1.20 \pm 0.01 \times 10^3 R_{\text{eff}}^{2.2 \pm 0.01}$ , suggesting that the clumps have nearly constant column density. This result is slightly steeper than the mass–size relationship of  $M_{\text{dust}} \propto R_{\text{eff}}^{1.67}$  found by Urquhart et al. (2014), based on masses derived for a sample of ATLASGAL clumps assuming a constant dust temperature of 20 K.

We compared the clump masses and sizes with the empirical threshold found by Kauffmann et al. (2010b) for clouds in which high-mass stars are usually found. While the clumps below this empirical threshold might or not form stars, we do expect that all the clumps above this threshold will eventually form at least one high-mass star. In this work we used a factor of 580 instead of the original 870 value. This is to account for the value of the opacity used in the column density calculation. The majority (78 per cent)

of the clumps in our sample lie above this limit, suggesting that they have sufficient mass and density to harbour or eventually form at least one high-mass star.

We also analysed the potential of clumps to evolve into open clusters. For this we assumed that the clumps have a 30 per cent star formation efficiency. Fig. 4 shows the parameter space in a plot mass versus radius allowed for open cluster progenitors (Portegies Zwart et al. 2010). We found that  $\sim 67$  per cent of the clumps in our sample have masses and sizes similar to those expected for the progenitors of open clusters.

The parameter space for clumps that have the mass and density of YMC progenitors is located in the upper end of this mass–size relationship. Five clumps within our sample lie in this region. Four of these five clumps have dust temperatures  $\geq 20$  K, and present



**Figure 4.** Mass–size relationship between the clumps. In this plot, the clumps are colour coded by dust temperature bins. Uncertainties in the measurements are shown with grey crosses. These uncertainties are dominated by the errors in the kinematic distance determination. The red dotted line shows the fit to the sample, and the black dotted lines show the ATLASGAL nominal sensitivity of  $0.25 \text{ mJy beam}^{-1}$ , which corresponds to  $2 \times 10^{21} \text{ cm}^{-2}$ . The yellow region highlights the parameter space for clouds that have been empirically shown to not harbour high-mass stars (Kauffmann et al. 2010a). Most of the clumps in our sample are above this limit, suggesting they will evolve into clusters harbouring high-mass stars. The red hatched region shows the parameter space for clump progenitors of open clusters (Portegies Zwart et al. 2010), assuming a star forming efficiency of 30 per cent. 69 per cent of the clumps lie within this region. The blue region highlights the parameter space for YMC progenitors (Bressert et al. 2012). There are just a few clumps meeting these criteria, which suggests that YMCs are rare in the plane of the Galaxy. Among the YMC progenitors one clump (marked with a black star in the plot) appears to be at the earliest stage of evolution.

some IR emission in the *Spitzer* maps, suggesting that star formation may have already begun within them. The small number of clumps that fulfil the requirements of YMC progenitors reinforces the idea that massive protoclusters are rare in the Galaxy.

In Table 2 we provide a summary of the physical properties of the clumps that have the characteristics of progenitors of clusters harbouring high-mass stars (HMPC), and clusters that may evolve into forming open clusters (OC) and YMCs.

### 3.4 Cold starless clumps

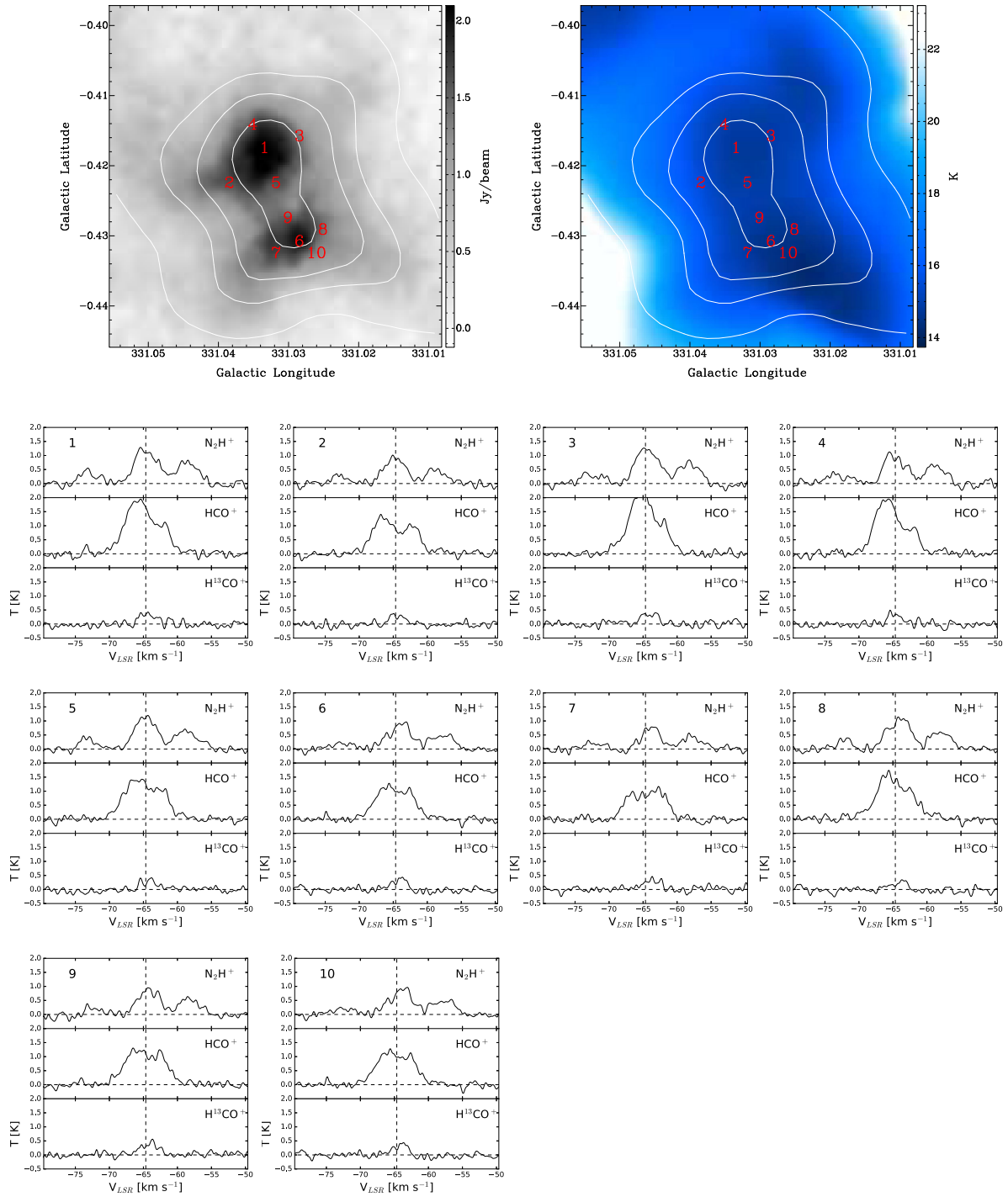
We found that 19 per cent of the clumps have dust temperatures  $\leq 20 \text{ K}$  and densities  $\geq 10^4 \text{ cm}^{-3}$  and thus have the properties of cluster-forming clumps in early evolutionary stages. These clumps are gravitationally bound: the median value of their virial parameter is  $\alpha \sim 0.8$ , suggesting they are likely to form stars. The mass–radius relationship for these cold clumps shows that the majority (88 per cent) have physical characteristics consistent with clump progenitors of high-mass stars. A large fraction (62 per cent) of the clumps are candidates for open cluster progenitors, with the majority of these open cluster progenitors potentially harbouring high-mass stars.

Within this sample of cold clumps, there are 24 clumps that lack emission in the  $70 \mu\text{m}$  *Herschel* maps [denominated as far-IRDC by Guzmán et al. (2015), for their lack of emission at far-IR wavelengths]. These clumps have a median dust temperature of

$14 \text{ K}$ , a median mass of  $\sim 10^3 M_\odot$ , and a median volume density of  $\sim 1 \times 10^4 \text{ cm}^{-3}$ . Consequently, they have the mass and density necessary to start forming stars. All but one of these clumps have physical properties consistent with clumps that will eventually form high-mass stars. Thus, this suggests that these cold, dark at  $70 \mu\text{m}$  clumps are in the very earliest stage of evolution, and thus they are excellent candidates to study the initial conditions of high-mass star formation.

We found only one cold clump, AGAL331.029-00.431\_S, with characteristics consistent with a YMC progenitor. Clump AGAL331.029-00.431\_S, composed of the MALT90 sources AGAL331.029-00.431\_S and AGAL331.034-00.419\_S, has a mean dust temperature of  $14 \text{ K}$ . The  $\text{H I}$  data towards this clump suggest that it is located at a far kinematic distance of  $10 \text{ kpc}$  (Whitaker et al., submitted). Wienen et al. (2015) located this clump at its near kinematic distance of  $\sim 4 \text{ kpc}$ . If we locate this clump at this distance, then, it would have a mass of  $6.6 \pm 0.8 \times 10^3 M_\odot$ , a density of  $3.9 \pm 0.3 \times 10^4 \text{ cm}^{-3}$ , and a dust bolometric luminosity of  $4.9 \pm 0.9 \times 10^3 L_\odot$ . However, this clump does not appear as an IR dark cloud in the *Spitzer* GLIMPSE images, in agreement with what is expected for a clump located at a distance of  $10 \text{ kpc}$ . At this distance the clump has a mass of  $4.1 \pm 0.2 \times 10^4 M_\odot$ , a density of  $7.9 \pm 0.3 \times 10^3 \text{ cm}^{-3}$  and a dust bolometric luminosity of  $3.1 \pm 0.6 \times 10^4 L_\odot$ .

The column density map shows a common envelope that surrounds both dust continuum peaks (AGAL331.029-00.431\_S and

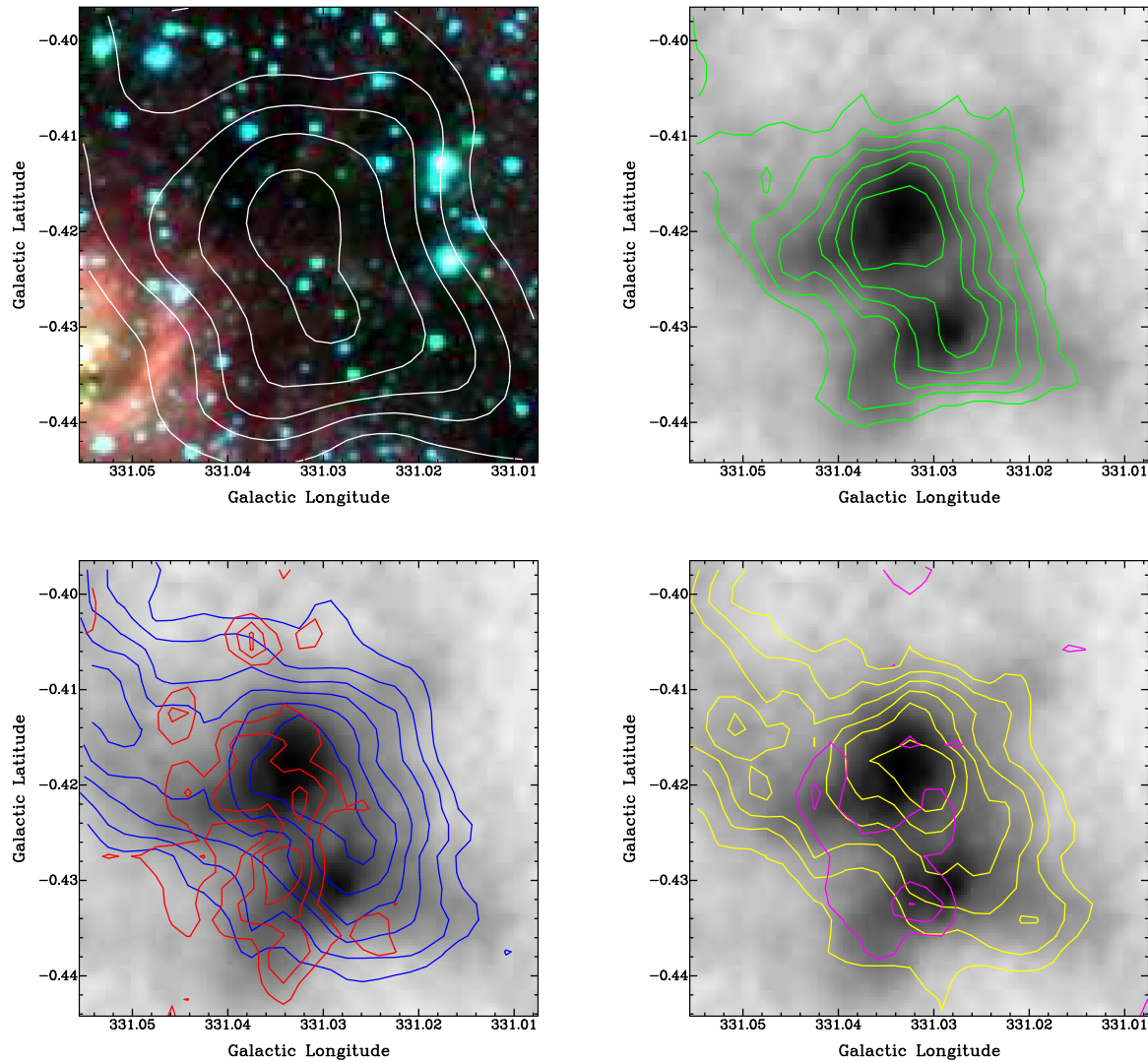


**Figure 5.** Dust continuum and molecular line emission towards the YMC progenitor candidate AGAL331.029-00.431\_S. Upper left-hand panel: ATLASGAL dust continuum emission overlaid with column density contours. Inner contours mark 70–90 per cent of the column density peak in this clump and the outer contour marks a column density of  $10^{22} \text{ cm}^{-2}$ . Upper right-hand panel: temperature of the clump overlaid with column density contours (values are the same as before). The levels of the column density are common to both dust peaks, indicating that it is a single coherent structure. The dust temperature map shows a low dust temperature across the clump, suggesting an early stage of evolution. The lower panels show spectra of  $\text{N}_2\text{H}^+$ ,  $\text{HCO}^+$  and  $\text{H}^{13}\text{CO}^+$  at the positions indicated in the images (red numbers). The positions are chosen to illustrate the molecular emission around and at the position of the two dust continuum peaks seen towards this clump. The  $\text{HCO}^+$  spectrum of the clump shows a blue peaked asymmetry in the whole clump, suggesting that it might be collapsing.

AGAL331.034-00.419\_S). The dust temperature map shows an almost homogeneous, cold dust temperature across the clump (Fig. 5). MALT90 observations show widespread emission of high-density tracers  $\text{N}_2\text{H}^+$ , HNC and  $\text{HCO}^+$  (1–0). The morphology exhibited

by these molecular lines is very similar, surrounding both dust continuum peaks, confirming that this is a physically coherent structure.

The  $\text{HCO}^+$  molecular line emission shows a blue peaked profile, which might indicate that this clump is collapsing. The blue peaked



**Figure 6.** IR, dust continuum and MALT90 molecular line emission towards AGAL331.029-00.431\_S. Upper left-hand panel: infrared colour image (red 8  $\mu\text{m}$ , green 4.5  $\mu\text{m}$  and blue 3.6  $\mu\text{m}$ ), overlaid with the column density determined towards this clump (as in Fig. 1). This cloud does not appear as an IR dark cloud, which is in agreement with the far kinematic distance determined via its molecular line emission. Upper right-hand panel:  $\text{N}_2\text{H}^+$  contours of the 40–90 per cent peak emission overlaid on the ATLASGAL dust continuum emission map. Lower left-hand panel:  $\text{HCO}^+$  contours of the 40–90 per cent peak emission in blue;  $\text{H}^{13}\text{CO}^+$  contours of the 70–90 per cent ( $5\sigma$  rms) of the peak emission in red. Lower right-hand panel:  $\text{HNC}$  contours of the 40–90 per cent peak emission in yellow;  $\text{HN}^{13}\text{C}$  contours of the 60–90 per cent ( $5\sigma$  rms) of the peak emission in magenta. This YMC candidate shows coherent emission in the molecular lines, while in the continuum two overdensities can be identified which correspond to the clumps identified in the ATLASGAL catalogue. The morphology is similar for  $\text{N}_2\text{H}^+$ ,  $\text{HNC}$  and  $\text{HCO}^+$ , with the peak of the molecular line emission located towards the upper dust overdensity seen in the ATLASGAL maps.  $\text{H}^{13}\text{CO}^+$  and  $\text{HN}^{13}\text{C}$  show different morphologies compared to the rest of the lines, with the peak of the emission between and closer to the lower dust overdensity, respectively. The overall morphology of the molecular emission exhibited by this clump, and its coherence in the  $(l, b, v)$  space, suggests that this clump has the mass and density to form a YMC.

profile is confirmed by the emission of the optically thin tracer  $\text{H}^{13}\text{CO}^+$  that has a peak at the velocity of the dip seen in  $\text{HCO}^+$  (Fig. 5).

The morphology exhibited by the optically thin tracers  $\text{H}^{13}\text{CO}^+$  and  $\text{HN}^{13}\text{C}$  (1–0) does not follow what is seen in the other molecules, with the peak of the integrated intensity located between the two overdensities seen in the dust continuum maps (Fig. 6). There is no emission from ‘hot core chemistry molecules’, such as  $\text{HC}^{13}\text{CCN}$ ,  $\text{HNCO}$ ,  $\text{CH}_3\text{CN}$  or  $\text{HC}_3\text{N}$  (1–0) that usually trace hot, dense gas that is heated by protostars, confirming the early stage of evolution of this clump. There is a marginal detection of the shock tracer  $\text{SiO}$ , which might indicate the presence of shocked gas, perhaps due to colliding gas inside this clump.

The apparent lack of any sign of current star formation, its mass and density, its coherence in the  $(l, b, v)$  space and the blue-peaked profile indicative of infall motions makes AGAL331.029-00.431\_S an excellent candidate of a YMC progenitor. Since the only YMC progenitor found is composed of two column density peaks, might argue that YMC progenitors are formed via the merging of two or more molecular clumps. We have to note that although our analysis places this clump at the far kinematic distance, further observations are needed to more accurately determine its distance and thus its true potential for forming a YMC. Nevertheless, the fact that we only see one, or possibly no, candidate for a YMC progenitor in this sample suggests that indeed YMCs are rare, or that the early stages of YMC formation are very quick. For example



Svoboda et al. (2016) estimate a lifetime of 0.03 Myr for YMC progenitors.

### 3.5 Volume-limited sample

We analysed the completeness of the MALT90 catalogue as a function of Galactic longitude to determine a volume-limited sample (VLS). The region where MALT90 is most complete, on which all the sources with fluxes  $>0.25$  Jy from the ATLASGAL catalogue were observed, lies between  $-50^\circ$  and  $-24^\circ$  in Galactic longitude. Within this region the MALT90 survey reached a sensitivity of  $200 M_\odot$  up to 5 kpc. There are 670 clumps in this volume, which corresponds to nearly half of the clumps used in this study. To check that we did not have any bias due to the incompleteness of the survey in certain regions of the Galactic plane, we reanalysed the data using only the VLS. We confirmed all trends of the VLS followed the same trends exhibited by the whole sample. Therefore, considering that the inclusion of sources from less complete regions of the survey does not seem to introduce important biases, we keep the full sample to have better statistics in the analysis.

## 4 CONCLUSIONS

We determined the physical properties (mass, virial mass, virial state, volume density and bolometric luminosity) of 1244 clumps located across the Galaxy. Analysing their molecular line emission and their dust column density, we determined that some column density peaks identified from the ATLASGAL maps are actually part of more extended coherent clumps.

Using the dust temperature of the clumps as proxy of their evolutionary stage, we determined how the physical properties of the clumps change as the clumps evolve. We found that the median mass of the clumps is  $\sim 10^3 M_\odot$  and shows no significant change with clump evolution. This suggests that we are sampling the same mass distribution in all evolutionary stages, and thus, that there is no mass bias in our sample.

The volume density monotonically decreases with dust temperature, while the virial parameter and bolometric luminosity monotonically increases with dust temperature. The changes in these parameters are consistent with the evolution of the clumps. At the initial stages the clumps are colder ( $T_{\text{dust}} < 20$  K), they are more bound ( $\alpha \sim 1$ ), their bolometric luminosity is lower ( $\sim 400 L_\odot$ ) and their median volume density is higher ( $\sim 10^4 \text{ cm}^{-3}$ ). This median volume density for the clumps in an early evolutionary stage is similar to the average star formation volume density of  $10^4 \text{ cm}^{-3}$  (Lada, Lombardi & Alves 2010), suggesting that these clumps are dense enough to form stars. As the clumps evolve and start forming stars their natal gas becomes hotter due to the radiation field of the newly forming stars, increasing their dust temperature ( $> 25$  K) and increasing their bolometric luminosity ( $\sim 2 \times 10^4 L_\odot$ ). At the same time, the clumps start to dissipate their surrounding gas, becoming more gravitationally unbound ( $\alpha > 1$ ) and decreasing their mean volume density ( $\sim 5 \times 10^3 \text{ cm}^{-3}$ ). This result suggests that dust temperature separation seems to be a better and more consistent indicator of the evolutionary stage of the clumps than classification based on their IR emission.

We also determined the potential of the clumps to evolve into stellar clusters harbouring high-mass stars or to form open clusters and YMCs. Most the clumps satisfy the criteria for high-mass protocluster candidates. Their range of dust temperatures suggest that they are in a wide range of evolutionary stages. We identified five clumps with characteristics of YMC candidates, which may form stellar clusters with masses  $\geq 10^4 M_\odot$ . The small number of YMCs

in our sample confirms that YMC progenitors in the Galactic plane are rare.

A significant fraction of the clumps (42 per cent) have low dust temperatures ( $< 20$  K) and therefore are good candidates for cluster-forming clumps in their earliest stage of evolution. Most of these cold clumps have physical properties consistent with their being progenitors of clusters that will harbour high-mass stars. A subsample of these clumps appear dark at  $70 \mu\text{m}$  (far-IRDC clumps; Guzmán et al. 2015). These clumps have very cold median dust temperatures (14 K), high masses ( $\sim 10^4 M_\odot$ ) and high volume densities ( $\sim 10^4 \text{ cm}^{-3}$ ); thus they likely represent the very earliest stages of high-mass star formation, and represent excellent targets to determine the initial conditions of high-mass star and cluster formation.

Only one clump has the physical characteristics of a YMC progenitor in a very early stage of evolution. It has a low dust temperature (14 K), high mass and density and no sign of evident current star formation at IR. So far, the only YMC progenitors known have been located in the harsh environment that surrounds the Galactic Centre ('The Brick'; Longmore et al. 2014; Rathborne et al. 2014a,b, 2015). The fact that we have only found one clump that can potentially evolve into a YMC reinforces the suggestion that YMC progenitors are rare, and that their earliest stages of formation are very short.

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## REFERENCES

- André P., Belloche A., Motte F., Peretto N., 2007, *A&A*, 472, 519
- Ao Y. et al., 2013, *A&A*, 550, A135
- Benjamin R. A. et al., 2003, *PASP*, 115, 953
- Bertoldi F., McKee C. F., 1992, *ApJ*, 395, 140
- Bressert E., Ginsburg A., Bally J., Battersby C., Longmore S., Testi L., 2012, *ApJ*, 758, L28
- Carey S. et al., 2005, Spitzer Proposal ID #20597
- Churchwell E. et al., 2009, *PASP*, 121, 213
- Contreras Y. et al., 2013a, *A&A*, 549, A45
- Contreras Y., Rathborne J., Garay G., 2013b, *MNRAS*, 433, 251
- Contreras Y., Garay G., Rathborne J. M., Sanhueza P., 2016, *MNRAS*, 456, 2041
- Copson E., 1965, *Asymptotic Expansions*. Cambridge Tracts in Mathematics. Cambridge Univ. Press, Cambridge
- Dunham M. K. et al., 2010, *ApJ*, 717, 1157
- Faúndez S., Bronfman L., Garay G., Chini R., Nyman L.-Å., May J., 2004, *A&A*, 426, 97
- Foster J. B. et al., 2011, *ApJS*, 197, 25
- Foster J. B. et al., 2013, *Publ. Astron. Soc. Aust.*, 30, 38
- Ginsburg A., Bressert E., Bally J., Battersby C., 2012, *ApJ*, 758, L29
- Green J. A., McClure-Griffiths N. M., 2011, *MNRAS*, 417, 2500
- Guzmán A. E., Sanhueza P., Contreras Y., Smith H. A., Jackson J. M., Hoq S., Rathborne J. M., 2015, *ApJ*, 815, 130



- Hill T., Burton M. G., Minier V., Thompson M. A., Walsh A. J., Hunt-Cunningham M., Garay G., 2005, *MNRAS*, 363, 405
- Hou S. et al., 2013, *ApJ*, 777, 157
- Jackson J. M. et al., 2013, *Publ. Astron. Soc. Aust.*, 30, 57
- Kauffmann J., Pillai T., Shetty R., Myers P. C., Goodman A. A., 2010a, *ApJ*, 712, 1137
- Kauffmann J., Pillai T., Shetty R., Myers P. C., Goodman A. A., 2010b, *ApJ*, 716, 433
- Kolpak M. A., Jackson J. M., Bania T. M., Clemens D. P., Dickey J. M., 2003, *ApJ*, 582, 756
- Lada C. J., Lombardi M., Alves J. F., 2010, *ApJ*, 724, 687
- Lada C. J., Forbrich J., Lombardi M., Alves J. F., 2012, *ApJ*, 745, 190
- Lis D. C., Carlstrom J. E., 1994, *ApJ*, 424, 189
- Lis D. C., Menten K. M., 1998, *ApJ*, 507, 794
- Lis D. C., Serabyn E., Zylka R., Li Y., 2001, *ApJ*, 550, 761
- Longmore S. N. et al., 2014, in Beuther H., Klessen R. S., Dullemond C. P., Henning T., eds, *Protostars and Planets VI*. Univ. Arizona Press, Tucson, AZ, p. 291
- McKee C. F., Ostriker E. C., 2007, *ARA&A*, 45, 565
- Molinari S. et al., 2010, *PASP*, 122, 314
- Morris M., Serabyn E., 1996, *ARA&A*, 34, 645
- Mueller K. E., Shirley Y. L., Evans N. J., II, Jacobson H. R., 2002, *ApJS*, 143, 469
- Nordhaus M. K. et al., 2008, in Frebel A., Maund J. R., Shen J., Siegel M. H., eds, *ASP Conf. Ser. Vol. 393*, New Horizons in Astronomy. Astron. Soc. Pac., San Francisco, p. 243
- Ormel C. W., Min M., Tielens A. G. G. M., Dominik C., Paszun D., 2011, *A&A*, 532, A43
- Portegies Zwart S. F., McMillan S. L. W., Gieles M., 2010, *ARA&A*, 48, 431
- Rathborne J. M. et al., 2014a, *ApJ*, 786, 140
- Rathborne J. M. et al., 2014b, *ApJ*, 795, L25
- Rathborne J. M. et al., 2015, *ApJ*, 802, 125
- Rathborne J. M. et al., 2016, *Publ. Astron. Soc. Aust.*, 33, e030
- Sanhueza P., Jackson J. M., Foster J. B., Garay G., Silva A., Finn S. C., 2012, *ApJ*, 756, 60
- Schuller F. et al., 2009, *A&A*, 504, 415
- Stephens I. W., Jackson J. M., Whitaker J. S., Contreras Y., Guzmán A. E., Sanhueza P., Foster J. B., Rathborne J. M., 2016, *ApJ*, 824, 29
- Svoboda B. E. et al., 2016, *ApJ*, 822, 59
- Tafalla M., Myers P. C., Caselli P., Walmsley C. M., 2004, *A&A*, 416, 191
- Tan J. C., Krumholz M. R., McKee C. F., 2006, *ApJ*, 641, L121
- Tan J. C., Beltrán M. T., Caselli P., Fontani F., Fuente A., Krumholz M. R., McKee C. F., Stolte A., 2014, in Beuther H., Klessen R. S., Dullemond C. P., Henning T., eds, *Protostars and Planets VI*. Univ. Arizona Press, Tucson, AZ, p. 149
- Urquhart J. S. et al., 2014, *A&A*, 568, A41
- Walmsley C. M., Guesten R., Angerhofer P., Churchwell E., Mundy L., 1986, *A&A*, 155, 129
- Wienen M. et al., 2015, *A&A*, 579, A91

## SUPPORTING INFORMATION

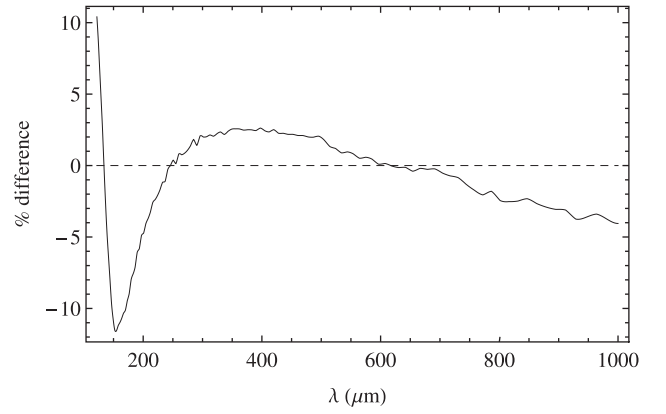
Supplementary data are available at [MNRAS](https://www.mnras.org/) online.

**Table 1.** Example of the summary table of physical properties derived for the MALT90 clumps with reliable kinematic distances.

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## APPENDIX A: DERIVATION OF THE BOLOMETRIC LUMINOSITY

The bolometric luminosity of the clumps is derived from extrapolating the emission seen in the continuum to shorter wavelengths.



**Figure A1.** Relative difference between equation (A1) and the dust absorption curve used in Guzmán et al. (2015).

To do this we assume that the dust absorption coefficient behaves as  $\nu^\beta$  with frequency, where  $\beta$  is the dust opacity spectral index. In the wavelength range used to derive the dust temperature and column density by Guzmán et al. (2015), the following is a good approximation for the mass absorption coefficient, assuming a gas-to-dust mass ratio of 100:

$$\kappa_\nu = 0.14 \left( \frac{250 \mu\text{m}}{\lambda_{\mu\text{m}}} \right)^\beta \text{ cm}^2 \text{ g}^{-1}, \quad (\text{A1})$$

where  $\lambda$  is the frequency of the emission. In equation (A1),  $\beta = 1.7$  gives a good fit. In Fig. A1 we show the percentage of difference between the approximation and the dust absorption coefficients from Ormel et al. (2011) used in Guzmán et al. (2015). Note that the approximation worsens towards wavelengths shorter than 120  $\mu\text{m}$ , thus equation (A1) is only valid for dust temperatures lower than  $\sim 25$  K. For higher dust temperatures, a large fraction of the radiation is emitted by dust associated with an absorption curve significantly different than that given by equation (A1).

The intensity received from a source of isothermal dust at temperature  $T$  is given by  $B_\nu(T)(1 - e^{-\tau_\nu})$ , where  $\tau_\nu = \tau_0(\nu/\nu_0)^\beta$  is the optical depth. The bolometric intensity is given by

$$I_{\text{bol}}(T) = \int_0^\infty B_\nu(T)(1 - e^{-\tau_\nu}) d\nu \quad (\text{A2})$$

$$= \frac{2k^4}{c^2 h^3} T^4 \int_0^\infty \frac{x^3}{e^x - 1} (1 - \exp(\tau_T x^\beta)) dx \quad (\text{A3})$$

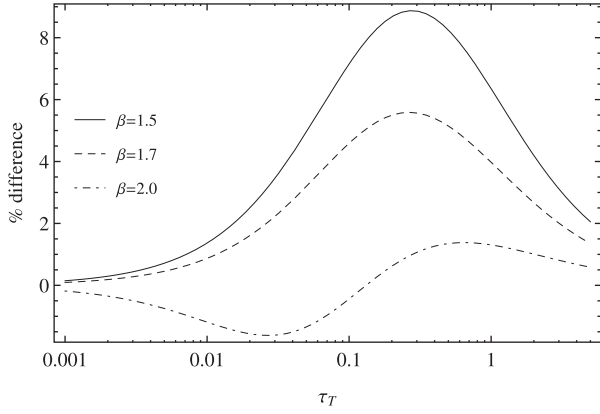
$$= \frac{15}{\pi^5} \sigma_{\text{SB}} T^4 \int_0^\infty \frac{x^3}{e^x - 1} (1 - \exp(\tau_T x^\beta)) dx, \quad (\text{A4})$$

where in equation (A3) we make the substitution  $x = h\nu/kT$ . In equations (A3) and (A4),  $\tau_T$  is the opacity at the frequency  $kT/h$  and  $\sigma_{\text{SB}}$  is the Stefan–Boltzmann’s constant.

Asymptotic forms of the integral in the right-hand side of equation (A4) for small and large  $\tau_T$  – optically thin and thick cases, respectively – are given by (e.g. Copson 1965, chapter 6)

$$\begin{cases} \tau_T \Gamma(\beta + 4) \zeta(\beta + 4) + O(\tau_T^2) & \tau_T \ll 1 \\ \frac{\pi^4}{15} + O(\tau_T^{-3/\beta}) & \tau_T \gg 1 \end{cases}, \quad (\text{A5})$$

where  $\Gamma$  and  $\zeta$  are the Gamma and Zeta functions, respectively and  $O$  is the big-O notation that indicates the asymptotic behaviour in the approximation’s residuals. In the optically thin case, the intensity (to first order) is proportional to  $\tau_T$ . In the optically thick case, on the other hand, the intensity reaches an asymptotic constant value.



**Figure A2.** Percentage of difference between the integral in equation (A4) and the approximation in equation (A6) for three values of  $\beta$ .

To interpolate between these two regimes, we use the following function:

$$\tau_T \Gamma(\beta + 4) \zeta(\beta + 4) \left( 1 + \frac{\tau_T \Gamma(\beta + 4) \zeta(\beta + 4)}{\pi^4/15} \right)^{-1}. \quad (\text{A6})$$

Equation (A6) has the same behaviour in the limit for small and large  $\tau_T$  as equation (A5), and therefore, as the integral in equation (A4). Fig. A2 shows the difference between the exact integral and the approximation for different values of  $\tau_T$  and  $\beta$ .

Finally, the bolometric luminosity of an isotropic source covering a solid angle  $\Omega_s$ , located at a distance  $D$ , is given by  $L_{\text{bol}} = I_{\text{bol}} \Omega_s D^2$ . Using the gas absorption coefficient  $\kappa_\nu$  given in equation (A1) (which already assumes a gas-to-dust ratio of 100) we can write the opacity  $\tau_T$  as  $N_{\text{gas}} \kappa_\nu$  with  $\nu = kT/h$ . Combining these expressions with equations (A4) and (A6), and changing to more convenient units, we write  $L_{\text{bol}}$  as

$$L_{\text{bol}} = 40 L_\odot \left( \frac{T}{10 \text{ K}} \right)^{5.7} \left( \frac{D}{\text{kpc}} \right)^2 \left( \frac{\Omega_s}{\text{arcmin}^2} \right) \times \left( \frac{N_{\text{gas}}}{\text{g cm}^{-2}} \right) C(T, N_{\text{gas}}), \quad (\text{A7})$$

where  $C(T, N_{\text{gas}})$  is the opacity and dust temperature correction factor (the rightmost term in equation A6):

$$C(T, N_{\text{gas}}) = (1 + 1.6 \times 10^{-3} N_{\text{gas}} T^{1.7})^{-1}. \quad (\text{A8})$$

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